The **STARK-B** database for spectral line broadening by collisions with charged particles

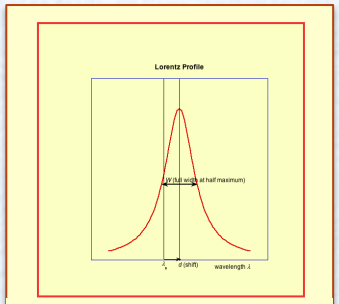
in the framework of the

European project **VAMDC**

(Virtual Atomic and Molecular Data Center)

S. Sahal-Bréchot

Observatoire de Paris, LERMA CNRS UMR 8112, France



http://stark-b.obspm.fr

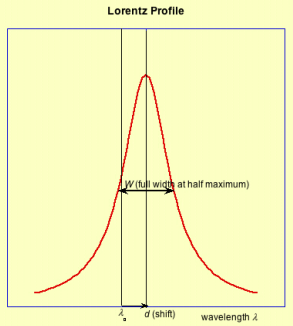
http://www.vamdc.org

STARK-B

Database for "Stark" widths and shifts of isolated spectral lines of atoms and ions, due to collisions with electrons and ions in a plasma

http://stark-b.obspm.fr

- **Calculated widths and shifts:** more than 150 pubs (*1984-2011*)
- **SCP theory updated** and **operated**
  by M.S. Dimitrijević and S. Sahal-Bréchot and colleagues
- **Technical development:**
  Paris Observatory, LERMA, by **N. Moreau, engineer**
  - *the database has been opened since September 2008: 99% of the existing data are currently implemented*
- It is a part of the atomic and molecular databases of the Paris Observatory
- Link (and mirror site in progress) to SerVO - Serbian Virtual Observatory
  http://servo.aob.rs/~darko/
- It is a part of VAMDC- Virtual Atomic and Molecular Data Centre
- it follows the standards of VAMDC and Virtual Observatories
  (Europe: IVOA *International Virtual Observatory Alliance*)



VAMDC Virtual Atomic and Molecular Data Centre
http://www.vamdc.org

IVOA
http://www.ivoa.net

STARK-B

widths and shifts of *isolated* spectral lines of atoms and ions, due to *collisions* with electrons and ions in a plasma

Overlapping lines and departures from impact approximation are outside the scope of STARK-B

Stark broadening of spectral lines can be applied to many subjects

- Astrophysics
- Laboratory plasmas
- Industrial plasmas

- Ionization degree ≥ 1%
- Moderately hot to very hot plasmas (5 10^3 to ~10^6 K)
- Moderate electron density (10^{13} to 10^{22} cm^{-3})

Diagnostics and Modelling in Astrophysics: Understanding of the evolution of stars

- *Thanks to considerable developments of*
  - *Ground based and space-born missions*
  - *Increased sensitivity (S/N) and spectral resolution*
  - *Powerful computers*

Interpretation of the faint observed spectrum :
faint objects, faint lines (trace elements)

- Line intensities + line profiles
- Continuum
- Spectroscopic diagnostics:
  - Temperatures
  - pressure
  - Abundances
  - Chemical stratification of the atmosphere

Modelling of atmospheres

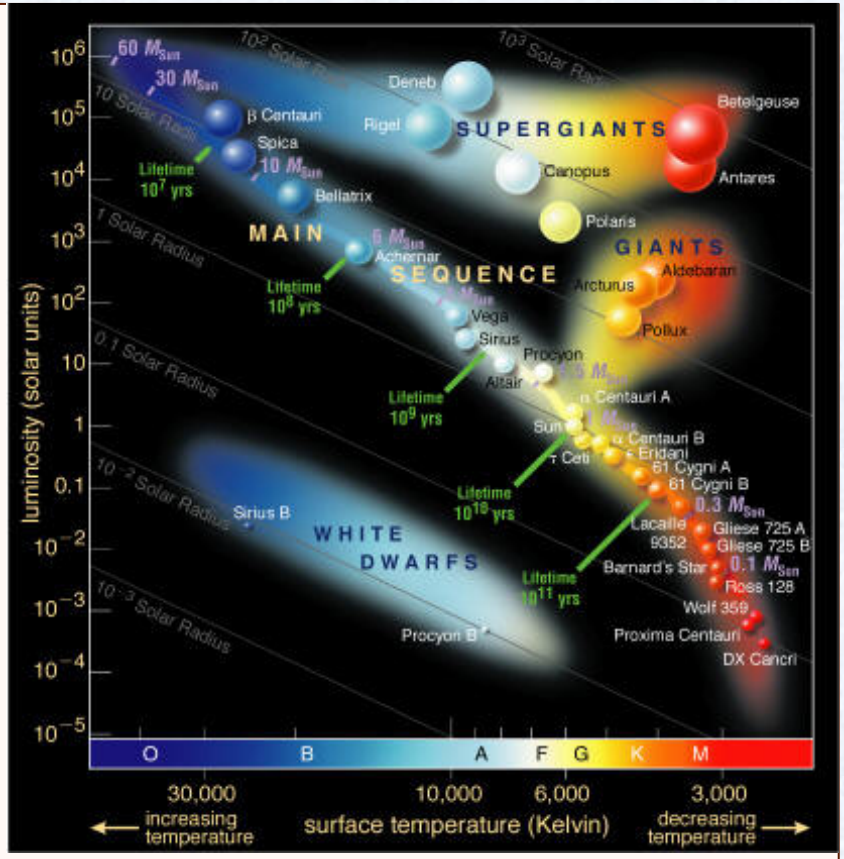
- Synthetic spectra: a great number of lines of a same element are required
- Radiative transfer

Stellar interiors studies and Asterosimology

- opacities: a great number of lines oh higly ionized elements are required
- Nuclear processes: formation of elements, chemical enrichment

Diagnostics and Modelling in Astrophysics: Understanding of the evolution of stars

- *Thanks to considerable developments of*
  - *Ground based and space-born missions*
  - *Increased sensitivity (S/N) and spectral resolution*
  - *Powerful computers*



Interpretation of the faint observed spectrum faint objects, faint lines (trace elements)

- Line intensities + line profiles
- Continuum
- Spectroscopic diagnostics:
  - Temperatures
  - pressure
  - Abundances
  - Chemical stratification of the atmosphere

Modelling of atmospheres

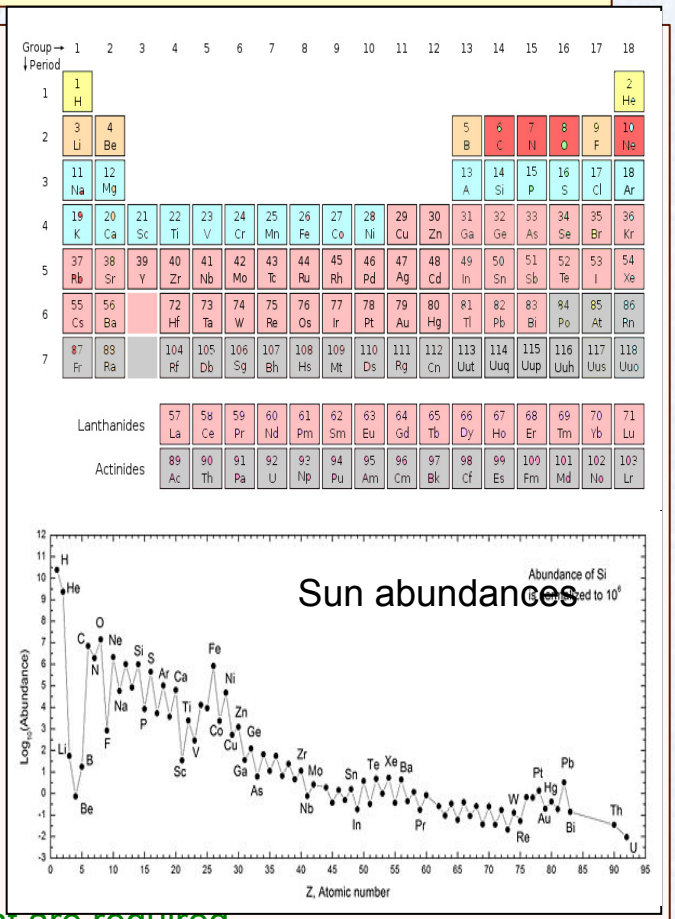
- Synthetic spectra: a great number of lines of a same element are required
- Radiative transfer

Stellar interiors studies and Asterosimology

- opacities: a great number of lines oh higly ionized elements are required
- Nuclear processes: formation of elements, chemical enrichment

Diagnostics and Modelling in Astrophysics: Understanding of the evolution of stars

- *Thanks to considerable developments of*
  - *Ground based and space-born missions*
  - *Increased sensitivity (S/N) and spectral resolution*
  - *Powerful computers*



Interpretation of the faint observed spectrum : faint objects, faint lines (trace elements)

- Line intensities + line profiles
- Continuum
- Spectroscopic diagnostics:
  - Temperatures
  - pressure
  - Abundances
  - Chemical stratification of the atmosphere

Modelling of atmospheres

- Synthetic spectra: a great number of lines of a same element are required
- Radiative transfer

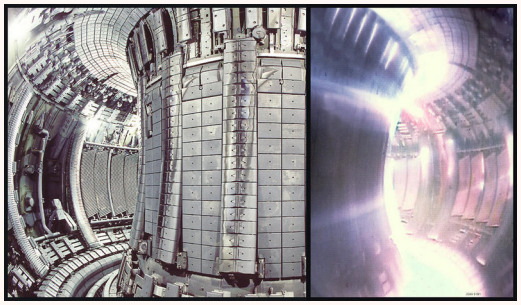
Stellar interiors studies and Asterosimology

- opacities: a great number of lines oh higly ionized elements are required
- Nuclear processes: formation of elements, chemical enrichment

Diagnostics and Modelling in Laboratory and Industrial plasmas

- *Thanks to developments and needs (devices an research)*
  - *Magnetic confinement fusion: moderately dense and hot plasmas (e.g. ITER)*
  - *Inertial confinement fusion (laser fusion, ion-beam fusion): dense and very hot plasmas*
  - *Low temperatures plasmas*
  - *Lighting discharges*

Analysis and interpretation of the spectrum in fusion devices and Tokamaks



- Light elements in the divertor and edge plasma regions
- Importance of Tin, Tungsten, and Aluminium
  (*Al-pellet Ablation Plasmas in Large Helical Device*)
  - Temperatures
  - Pressure

Progress in low-energy light sources

- Discharge lamps and lighting :
  - optimisation of performances
  - cold light from hot atoms and molecules
    - (white light 3000-5000K, discharge up to 45000K),
    - high electron density (strong and broad emission in the visible spectrum
  - Fluorescent lamps: improving efficacity (phosphors)
- Rare earth elements Dy, Ho, Ce: excellent radiation sources
- HID (*High Intensity Discharge*): MH (*Metal Halide*) lamps, e.g. Dy I_3, In I, ZnI3
- (LED light-emitting diodes)

STARK-B data: Basic approximations (theory and calculations)

- *Impact approximation*
  - *Collisions between radiators and perturbers act independently and are additive*

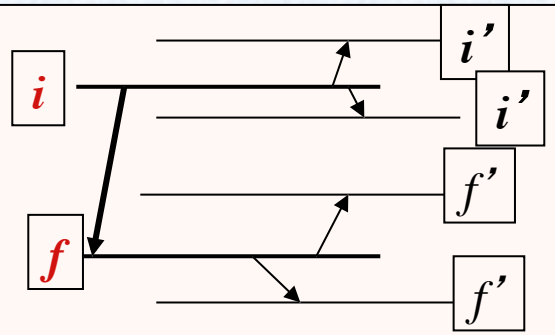
$\rho_{typ}/v \ll \Delta T$ i.e. $\rho_{typ} \ll N^{-1/3}$

- *Complete collision approximation*
  - ➔ *line broadening theory becomes an application of the theory of collisions*

The atom has no time to emit or absorb a photon during the collision process, *the collision is not broken off.*

- *Isolated lines*
  - *Neighbouring levels do not overlap*

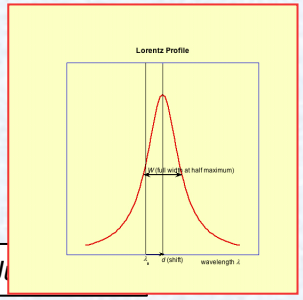
➔ *Lorentz profile; S-matrix, cross-sections*



$$W = N \int v\, f(v) \left(\sum_{i' \neq i} \sigma_{ii'}(v) + \sum_{f' \neq f} \sigma_{ff'}(v) + \sigma_{el}(v) \right)$$

- *LS coupling: fine (hyperfine) structure can be neglected during the collision (S or I : no time to rotate)*
  - ➔*The fine structure (hyperfine) components have the same width and the same shift, thet of the multiplet*

- *High densities:*
  - ➔*Debye screening effect*



STARK-B
Methods of calculations of the data
1. The scattering S-matrix

- **SCP: Semi-Classical** : atom = quantum description (atomic structure), perturber= particle moving on a classical path

  + **Perturbation theory (2nd order)**

  – (Sahal-Bréchot, A&A1970 and further papers, 6-8 basic papers)
    - unitarity and symmetrization of the S-matrix, adequate cut-offs
    - hyperbolae for ion-electron and ion-ion (1970),
    - complex atoms (1974), very complex (2008)
    - Feshbach resonances for ion-electron collisions (1977)
    - Updated and operated with MS. Dimitrijević (1984 and after)
    - accuracy: 20%, sometimes better, sometimes worse

- **MSE: Modified Semi-Empirical** : atom = simplified quantum description
  Dimitrijević and colleagues, JQSRT1980 and further A&A papers)

STARK-B
Methods of calculations of the data
2 - The atomic structure

- NIST atomic data
- VALD (Vienna Atomic Line Database)
- Coulomb approximation with quantum defect *(Bates & Damgaard 1949)*
- TOPbase : R-matrix in LS coupling
- Cowan code : HFS multi-conf with exchange and relativistic effects (by perturbations)
- SUPERSTRUCTURE : scaled Thomas-Fermi-Dirac-Amaldi potential + relativistic effects (Breit-Pauli)

STARK-B
Methods of calculations of the data
3. Calculations leading to a great number of data

Atomic structure coupled to the S-matrix calculation:

Many widths and shifts for a set of several temperatures and densities in a same run

- Ab initio calculations: no external data insertion
  - SST + SCP: data for 100 lines and more in a same run
  - Cowan code + SCP : in progress (case of Pb IV, more than 100 lines in a same run)
- Coupling to atomic structure databases:
  - TopBase : data for 150 lines in a same run
  - VALD: data for more than 1000 lines in a same run (case of CII for white dwarfs)

STARK-B:
Elements (atoms and ions) currently inserted

Ag I,
Al I, Al III, Al XI
Ar I, Ar II, Ar II, Ar VIII
Au I
B II, B III
Ba I, Ba II
Be I, Be II, Be III
Br I
C II, C III, C IV, C V
Ca I, Ca II, Ca V, Ca IX, Ca X
Cd I, Cd II
Cl I, Cl VII
Cr I, Cr II
Cu I
F I, F II, F III ,F IV, F V, F VI, F VII
Fe II
Ga I
Ge I
He I
Hg II
I I
In II, In III
K I, K VIII, K IX
Kr I, Kr II, Kr VIII

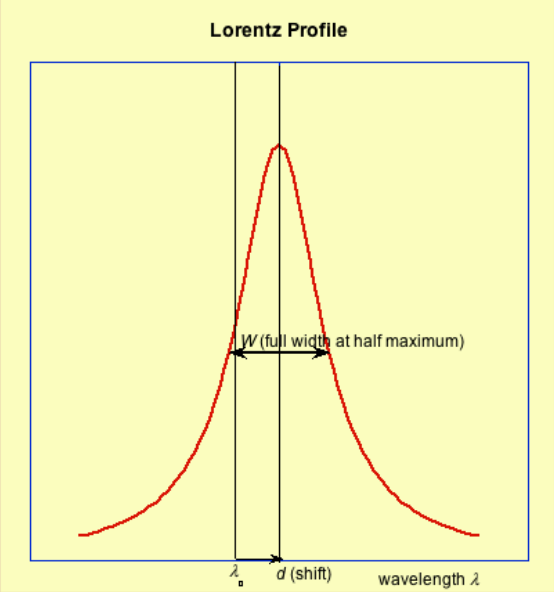
Li I, Li II
Mg I, Mg II, Mg XI
Mn II
N I, N II, N III, N IV, N V
Na I, Na X
Ne I, Ne II, Ne II, Ne III, Ne IV, Ne V, Ne VIII
Ni II
O I, O II, *O III (in progress)*, O IV, O V, O VI; O VII
P IV, P V
Pb IV
Pd I
Rb I
S III, S IV, S V, S VI
Sc III, Sc X, Sc XI
Se I
Si I, Si II, Si IV, Si V, Si VI, Si XI, Si XII, Si XIII
Sr I
Te I
Ti IV, Ti XII, Ti XIII
Tl III
V V, V XIII
Y III
Zn I
C II in progress (150 lines, 4 temperatures, 3 densities)

STARK-B http://stark-b.obspm.fr

HOMEPAGE | INTRODUCTION | DATA DESCRIPTION | ACCESS TO THE DATA | UPDATES | CONTACT

STARK-B

Database for "Stark" broadening of isolated lines of atoms and ions in the impact approximation



The STARK-B database is now fully opened though not yet complete.

Last data update : 2012-03-30

MOLAT | l'Observatoire de Paris | Astronomical Observatory, Belgrade

Introduction

This is a database of calculated widths and shifts of isolated lines of atoms and ions due to electron and ion collisions.

This database is devoted to modellisation and spectroscopic diagnostics of stellar atmospheres and envelopes. In addition, it is also devoted to laboratory plasmas, laser equipments and technological plasmas. So, the domain of temperatures and densities covered by the tables is wide and depends on the ionization degree of the considered ion. The temperature can vary from several thousands for neutral atoms to several hundred thousands of Kelvin for highly charged ions. The electron or ion density can vary from 10^{12} (case of stellar atmospheres) to several 10^{19} cm^{-3} (some white dwarfs and some laboratory plasmas).

The impact approximation and the isolated line approximation are applied, so that the line profile is Lorentzian. The basis for calculations is the computer code which evaluates electron and ion impact broadening of isolated spectral lines of neutral atoms and ions, using the semiclassical-perturbation approach developed by Sahal-Bréchot (1969ab, 1974), and supplemented in Fleurier etal. (1977), see below. This computer code has been updated by Dimitrijevic and Sahal-Bréchot in their series of papers, Dimitrijevic and Sahal-Bréchot (1984) and following papers. The data are derived from this series of papers and are cited in the tables.

The accuracy of the data varies from about 15-20 percent to 35 percent, depending on the degree of excitation of the upper level, and on the quality of the used atomic structure entering the calculation of scattering S-matrix leading to the widths and shifts. The more the upper level is excited, the more the accuracy is good. In the earlier papers, the used atomic structure was the so-called "Bates and Damgaard" one (Coulomb wavefunctions + quantum defect). More recent atomic structure data are introduced in the latest papers. The reader is invited to refer to the papers cited in the tables for the used atomic data and atomic levels, and for more details.

The impact approximation

The impact approximation is valid when the mean duration τ of a collision is much smaller than the mean interval ΔT between two collisions (Baranger 1958 abc). ΔT is of the order of the inverse of the colisional line width γ expressed in angular frequency units.

τ can be written as :

$\tau = \langle\rho\rangle / \langle v\rangle$,

where $\langle\rho\rangle$ is a typical impact parameter and $\langle v\rangle$ the mean velocity of the collider :

$\langle v\rangle = (8kT / \pi\mu)^{1/2}$,

μ being the reduced mass, T the temperature, and k the Bolzmann constant.

An order of magnitude of $\langle\rho\rangle$ can be derived from the line width γ and is obtained by writing (N being the density of the perturbers)

$\gamma = N\langle v\rangle\pi\langle\rho\rangle^2$.

The validity condition of the impact approximation can be written as

$N V \ll 1$,

$V = \pi\langle\rho\rangle^3$ is the collision volume (Baranger 1958abc).

The impact values of the widths and shifts are given in the tables, except when $N V > 0.5$. Then the cells are empty and marked by an asterisk preceding the cell. Widths values for $0.1 < N V \leq 0.5$ are marked by an asterisk in the cell preceding the value. See the Section "Data description" for more details.

In the far wings, $\Delta\omega$ being the detuning in angular frequency units, the validity condition for the generalized impact approximation becomes

$\tau\Delta\omega \ll 1$.

When the impact approximation is not valid (especially for ion colliders), the quasistatic approximation can be used. As shown by Baranger (1962) for ion emitters and polarization interaction potenial, and by Sahal-Bréchot (1991) for the quadrupolar interaction which is in fact dominant due to the Coulomb repulsion, the quasistatic broadening is completely negligible in the wings. For neutrals emitters, the polarization part of the interaction is most often dominant and can be obtained by the A parameter of Griem (1974). This A parameter is provided in a few tables, where it is calculated with the method described by by Ben Nessib et al. (1996).

The isolated line approximation

When the impact approximation is not valid (especially for ion colliders), the quasistatic approximation can be used. As shown by Baranger (1962) for ion emitters and polarization interaction potenial, and by Sahal-Bréchot (1991) for the quadrupolar interaction which is in fact dominant due to the Coulomb repulsion, the quasistatic broadening is completely negligible in the wings. For neutrals emitters, the polarization part of the interaction is most often dominant and can be obtained by the A parameter of Griem (1974). This A parameter is provided in a few tables, where it is calculated with the method described by by Ben Nessib et al. (1996).

The isolated line approximation

At high densities or for lines arising from high levels, the electron impact width becomes comparable to the separation $\Delta E(nl, nl\pm1)$ between the perturbing energy levels and the initial or final level : the corresponding levels become degenerate and the isolated line approximation is invalid (Griem 1974). In order to check the validity of this approximation, we have defined a parameter C in Dimitrijevic and Sahal-Bréchot (1984) which is given in the tables. See "Data description" for more details.

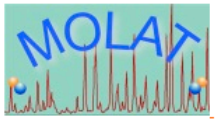
The semiclassical perturbation approximation (SCP)

When the impact approximation is valid, the collisional broadening becomes an application of the theory of collisions (Baranger 1958abc). In the semiclassical approximation, rectilinear trajectories are used for neutral emitters (or absorbers), and hyperbolic trajectories for ionic emitters (or absorbers) colliding with charged particles. Within the second order perturbation approximation, dipolar, polarization and quadrupolar interactions are taken into account (Sahal-Bréchot 1969ab and earlier papers), updated for complex atoms and ions (Sahal-Bréchot 1974). The details of calculations of the widths and shifts can be found on these papers. For ionic emitters, the original computer code has been updated by including Feshbach resonances in elastic and fine structure transitions by using the semiclassical limit of the Gailitis formula (Fleurier et al. 1977). Debye shielding effect is also taken into account. It is negligible at low densities or for lines arising from low levels. Then widths and shifts are proportional to the density.

Other data

When SCP calculations cannot be performed, this database will include data obtained using the MSE (Modified Semi-Empirical) method (Dimitrijevic & Konjevic 1980, Dimitrijevic 1982, Dimitrijevic & Krsljanin 1986), supplemented by Popovic & Dimitrijevic (1996) for complex atoms. These data will be included in a future version.
selected experimental data will also be included in another future version.

- Baranger, M. 1958a, Phys. Rev. A, 111, 481
- Baranger, M. 1958b, Phys. Rev. A, 111, 494
- Baranger, M. 1958c, Phys. Rev. A, 112, 855
- Baranger M., 1962 "Spectral line broadening by plasmas", edited by D. R. Bates. Library of Congress Catalog Card Number 62-13122. in "Atomic and Molecular Processes", pp. 493-548, Acad. Press Inc., New-York
- Ben Nessib, N., Ben Lakhdar, Z. et Sahal-Bréchot, S., 1996,, Phys. Scr., 54, 608-613.
- Dimitrijevic, M.S., and Sahal-Bréchot, S.: 1984,JQSRT 31, 301-313
- Dimitrijevic M.S., & Konjevic J., 1980, JQSRT, 24, 451
- Dimitrijevic M.S, 1982, A&A, 112, 251
- Dimitrijevic M.S. & Krsljanin, V. 1986, A&A, 165, 269
- Fleurier C., Sahal-Bréchot, S., and Chapelle, J.: 1977, JQSRT, 17, 595-604
- Griem, H. R. 1974, "Spectral line broadening by plasmas", Pure and Applied Physics, New York: Academic Press, USA
- Popovic L. C., & Dimitrijevic M.S., 1996, Phys. Scr., 53, 325
- Sahal-Bréchot, S.: 1969a, A&A 1, 91-123
- Sahal-Bréchot, S.: 1969b, A&A 2, 322-354
- Sahal-Bréchot, S.: 1974, A&A 35, 319-321
- Sahal-Bréchot, S.: 1991, Astron.Astrophys. 245, 322-330



Data description

Periodic table of elements

Click on a yellow case corresponding to the chosen element and then on an ionization degree. There are no data on the non-coloured cases. Then a new page appears, requiring your detailed choice.

Tables description

Column 1

Perturber density N in cm^{-3}

Column 2

Lower level or lower term

Column 3

Upper level or upper term

> #### Comments to columns 2 and 3
>
> When the fine structure splitting is small, namely if the difference between energy levels of a same multiplet is small compared to the distance to the next level linked by an allowed transition, all the fine structure lines of a same multiplet have the same width and shift. In that case the data are given for the multiplet only and for an average wavelength for the whole multiplet. If needed, the width value for a particular line within a multiplet can be obtained from :
>
> $W_{line} = W_{mult}\, I2_{line} / \lambda^2_{mult}$
>
> Idem for the shift

Column 4

Multiplet when it is available

> #### NB
>
> It is the multiplet number generated on line from the NIST Atomic Spectra Database *. Therefore we have chosen not to select any wavelength range, because the multiplets numbers vary if the the selected wavelength range varies. In addition, the data for multiplets as a whole are only generated on the NIST Atomic Spectra Database if all fine structure components are known and if it is in LS coupling. * NIST Atomic Spectra Database (version 3.1.5), [Online].
>
> Available: http://physics.nist.gov/asd3 Ralchenko, Yu., Kramida, A.E., Reader, J., and NIST ASD Team (2008), National Institute of Standards and Technology, Gaithersburg, MD.

Column 5

Wavelength in Å

> #### Comment to column 5 :
>
> These wavelengths are calculated wavelengths with the computer code. In particular, they are averaged over the multiplet when multiplet data are given.

Column 6

Parameter C for the validity condition of the isolated line approximation

> #### Comment to column 6
>
> The isolated line approximation is valid for a kind of perturbers a (a = electrons, protons, He II, ...) if C/W_a is higher than the corresponding perturber density. For a perturber density N lower than N_l (cm^{-3})=C/W_a,the line can be treated as isolated even if a weak forbidden component due to the failure of this approximation remains in the wing. W_a is the full width at half-intensity given in the

Column 6
Parameter C for the validity condition of the isolated line approximation

> Comment to column 6
> The isolated line approximation is valid for a kind of perturbers a (a = electrons, protons, He II, ...) if C/W_a is higher than the corresponding perturber density. For a perturber density N lower than N_l (cm^{-3})=C/W_a,the line can be treated as isolated even if a weak forbidden component due to the failure of this approximation remains in the wing. W_a is the full width at half-intensity given in the coresponding following columns (9, 11, 13...). See the "Introduction" for definition of the validity condition of the isolated line approximation.

Column 7
Temperature T in Kelvin

Column 8
A (quasistatic parameter for neutral atoms, cf. Introduction for details) if available

Column 9
Full width at half intensity We in Å (electron colliders)

Column 10
Shift de in Å (electron colliders). A positive shift is towards the red, a negative one is towards the blue

- Empty cells which are not preceded by an asterisk mean that the data are not available
- Empty cells which are preceded by an asterisk mean that the impact approximation is not valid, because NV > 0.5 (cf. Introduction for details), and thus the coresponding data are not provided
- Non-empty cells preceded by an asterisk mean that the impact approximation reachs its limit of validity, 0.1 < NV ≤ 0.5 (cf. Introduction for details) (cf. Introduction for details)

> NB
> When the shift is negative, due to the additional minus sign, only the width value is marked with the asterisk.

Columns 11 and 12
same as columns 9 and 10, but for protons colliders (subscript p)

Columns 13 and following columns :
same as columns 9 and 10, but for other ion colliders (other corresponding subscripts)

> NB
> Some widths and shifts appear at medium and not at low densities. This means that they are proportional with the density. Thus data at low densities can be deduced from those at medium densities by linear interpolation with the perturber density.

MOLAT — l'Observatoire de Paris — Astronomical Observatory, Belgrade

HOMEPAGE | INTRODUCTION | DATA DESCRIPTION | ACCESS TO THE DATA | UPDATES | CONTACT

Choose an element and an ionization degree

H																	He
Li	Be											B	C	N	O	F	Ne
Na	Mg											Al	Si	P	S	Cl	Ar
K	Ca	Sc	Ti	V	Cr	Mn	Fe	Co	Ni	Cu	Zn	Ga	Ge	As	Se	Br	Kr
Rb	Sr	Y	Zr	Nb	Mo	Tc	Ru	Rh	Pd	Ag	Cd	In	Sn	Sb	Te	I	Xe
Cs	Ba	La	Hf	Ta	W	Re	Os	Ir	Pt	Au	Hg	Tl	Pb	Bi	Po	At	Rn
Fr	Ra	Ac															
				Ce	Pr	Nd	Pm	Sm	Eu	Gd	Tb	Dy	Ho	Er	Tm	Yb	Lu
				Th	Pa	U	Np	Pu	Am	Cm	Bk	Cf	Es	Fm	Md	No	Lr

MOLAT

l'Observatoire de Paris

HOMEPAGE | INTRODUCTION | DATA DESCRIPTION | ACCESS TO THE DATA | UPDATES | CONTACT

Choose an element and a ionization degree

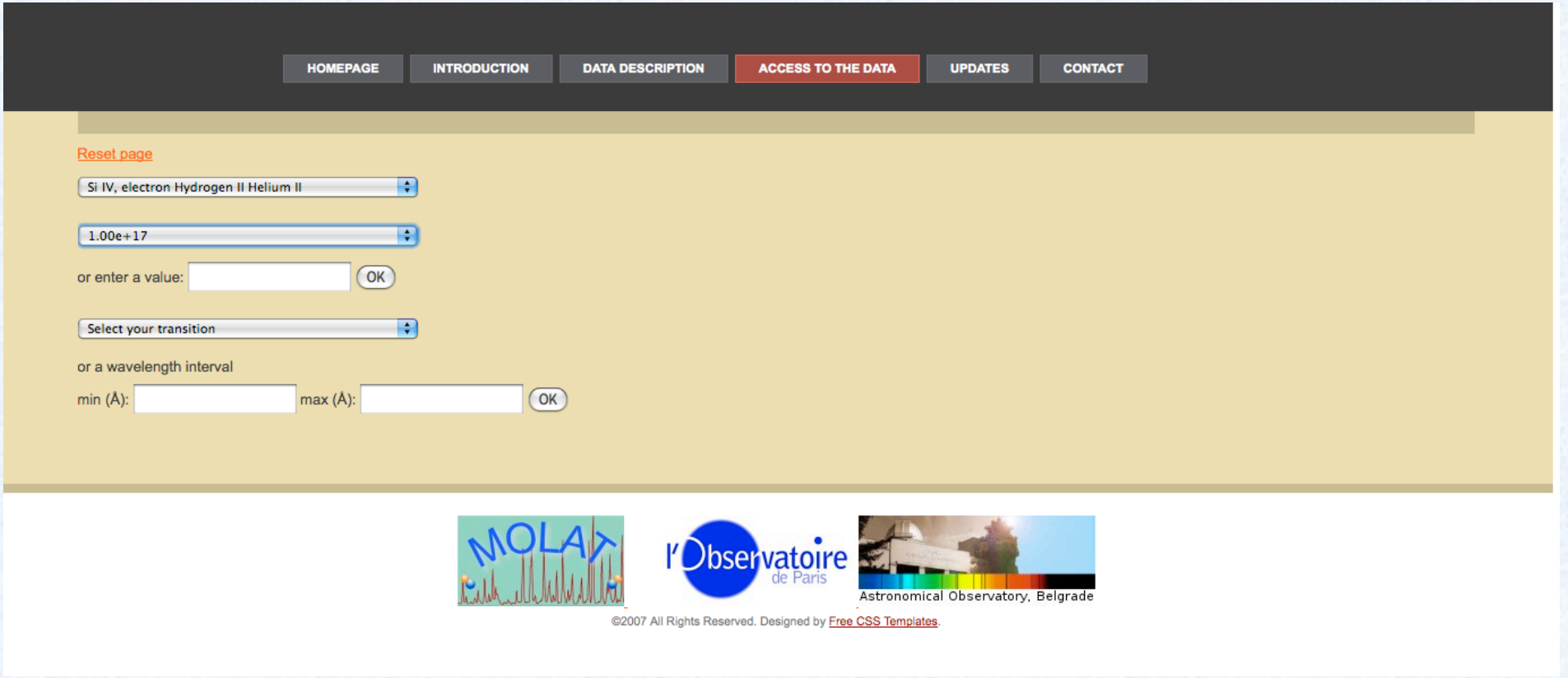
- Si I
- Si IV
- Si V
- Si VI
- Si XI
- Si XII
- Si XIII

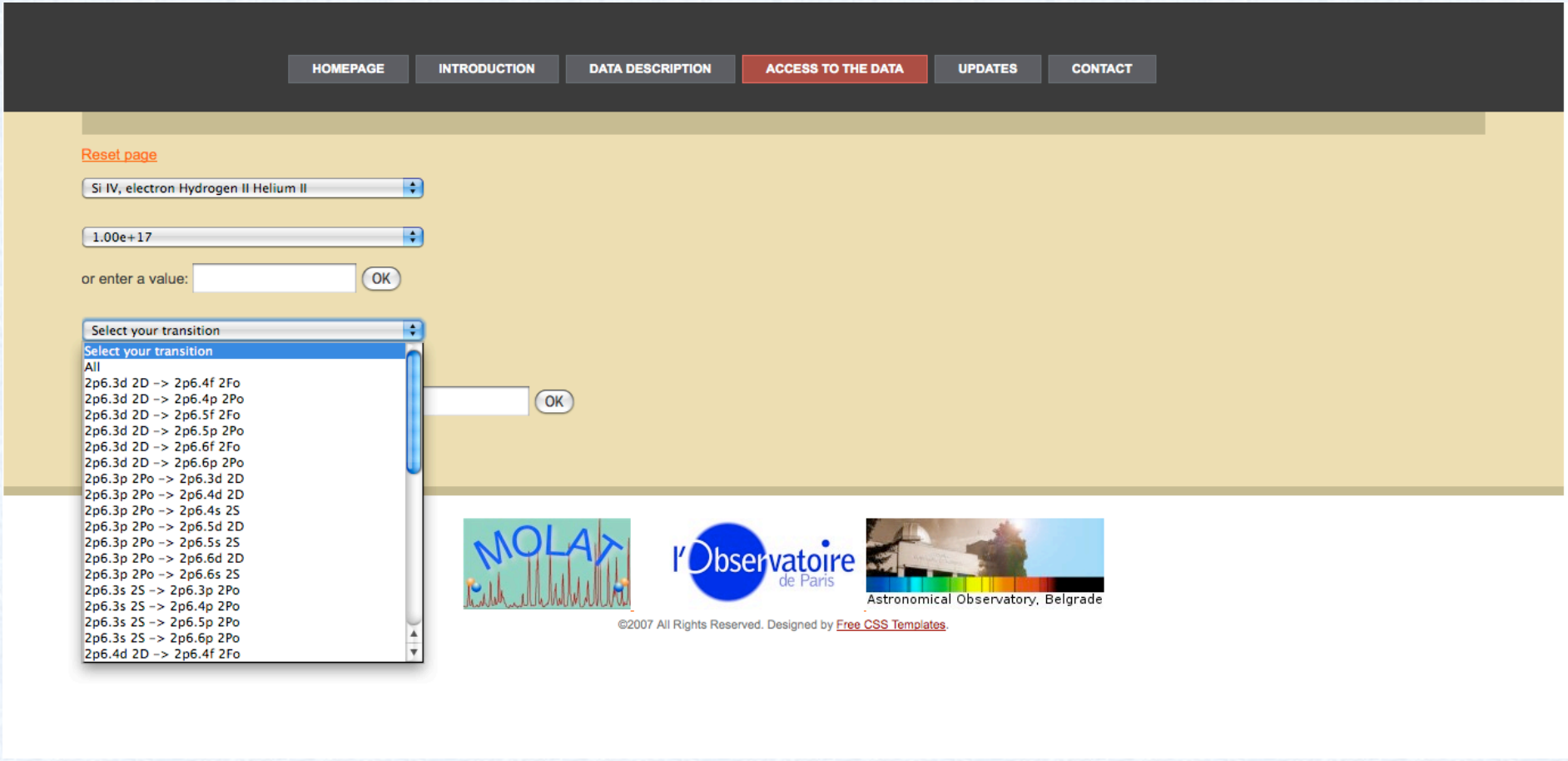
H																	He
Li	Be											B	C	N	O	F	Ne
Na	Mg											Al	Si	P	S	Cl	Ar
K	Ca	Sc	Ti	V	Cr	Mn	Fe	Co	Ni	Cu	Zn	Ga	Ge	As	Se	Br	Kr
Rb	Sr	Y	Zr	Nb	Mo	Tc	Ru	Rh	Pd	Ag	Cd	In	Sn	Sb	Te	I	Xe
Cs	Ba	La	Hf	Ta	W	Re	Os	Ir	Pt	Au	Hg	Tl	Pb	Bi	Po	At	Rn
Fr	Ra	Ac															
			Ce	Pr	Nd	Pm	Sm	Eu	Gd	Tb	Dy	Ho	Er	Tm	Yb	Lu	
			Th	Pa	U	Np	Pu	Am	Cm	Bk	Cf	Es	Fm	Md	No	Lr	

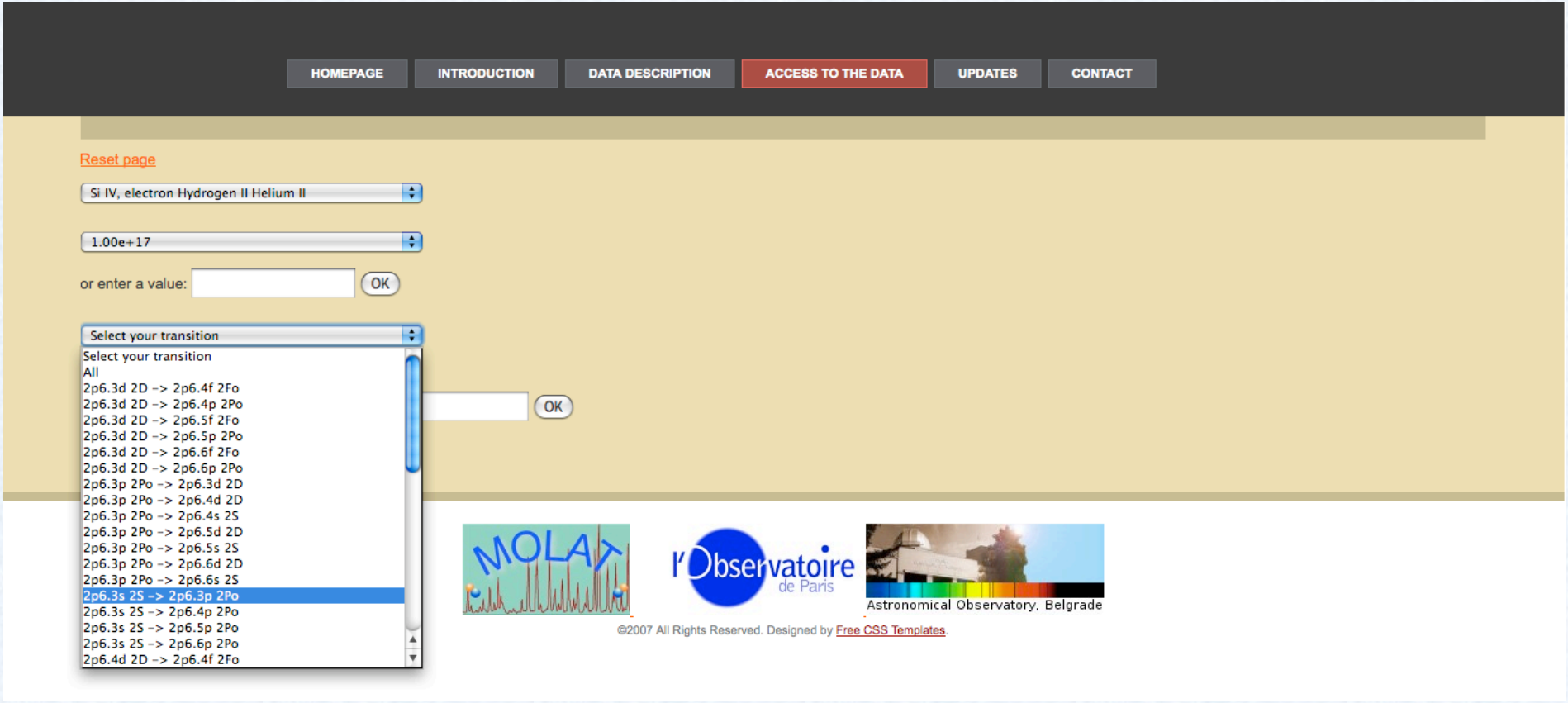
MOLAT | l'Observatoire de Paris | Astronomical Observatory, Belgrade


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Select your dataset
Select your dataset
Si IV, electron Hydrogen II Helium II
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Si IV, electron Hydrogen II Helium II
Select your Perturber density (cm-3)
or enter a value:
OK
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Si IV, electron Hydrogen II Helium II
1.00e+17
or enter a value:
OK
Select your transition
or a wavelength interval
min (Å):
max (Å):
OK
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Si IV, electron Hydrogen II Helium II
1.00e+17
or enter a value:
OK
Select your transition
Select your transition
All
2p6.3d 2D -> 2p6.4f 2Fo
2p6.3d 2D -> 2p6.4p 2Po
2p6.3d 2D -> 2p6.5f 2Fo
2p6.3d 2D -> 2p6.5p 2Po
2p6.3d 2D -> 2p6.6f 2Fo
2p6.3d 2D -> 2p6.6p 2Po
2p6.3p 2Po -> 2p6.3d 2D
2p6.3p 2Po -> 2p6.4d 2D
2p6.3p 2Po -> 2p6.4s 2S
2p6.3p 2Po -> 2p6.5d 2D
2p6.3p 2Po -> 2p6.5s 2S
2p6.3p 2Po -> 2p6.6d 2D
2p6.3p 2Po -> 2p6.6s 2S
2p6.3s 2S -> 2p6.3p 2Po
2p6.3s 2S -> 2p6.4p 2Po
2p6.3s 2S -> 2p6.5p 2Po
2p6.3s 2S -> 2p6.6p 2Po
2p6.4d 2D -> 2p6.4f 2Fo
OK
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Si IV, electron Hydrogen II Helium II
1.00e+17
or enter a value:
OK
Select your transition
Select your transition
All
2p6.3d 2D -> 2p6.4f 2Fo
2p6.3d 2D -> 2p6.4p 2Po
2p6.3d 2D -> 2p6.5f 2Fo
2p6.3d 2D -> 2p6.5p 2Po
2p6.3d 2D -> 2p6.6f 2Fo
2p6.3d 2D -> 2p6.6p 2Po
2p6.3p 2Po -> 2p6.3d 2D
2p6.3p 2Po -> 2p6.4d 2D
2p6.3p 2Po -> 2p6.4s 2S
2p6.3p 2Po -> 2p6.5d 2D
2p6.3p 2Po -> 2p6.5s 2S
2p6.3p 2Po -> 2p6.6d 2D
2p6.3p 2Po -> 2p6.6s 2S
2p6.3s 2S -> 2p6.3p 2Po
2p6.3s 2S -> 2p6.4p 2Po
2p6.3s 2S -> 2p6.5p 2Po
2p6.3s 2S -> 2p6.6p 2Po
2p6.4d 2D -> 2p6.4f 2Fo
OK
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Si IV, electron Hydrogen II Helium II
1.00e+17
or enter a value:
OK
2p6.3s 2S -> 2p6.3p 2Po
or a wavelength interval
min (Å):
max (Å):
OK
Select your temperature
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.


HOMEPAGE
INTRODUCTION
DATA DESCRIPTION
ACCESS TO THE DATA
UPDATES
CONTACT
Reset page
Si IV, electron Hydrogen II Helium II
1.00e+17
or enter a value:
OK
2p6.3s 2S -> 2p6.3p 2Po
or a wavelength interval
min (Å):
max (Å):
OK
Select your temperature
Select your temperature
All
20000
50000
80000
100000
150000
200000
MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade
©2007 All Rights Reserved. Designed by Free CSS Templates.

Reset page

Si IV, electron Hydrogen II Helium II

1.00e+17

or enter a value: OK

2p6.3s 2S -> 2p6.3p 2Po

or a wavelength interval

min (Å): max (Å): OK

All

When using these data, please refer to the original papers and to this database as :

Sahal-Bréchot, S., Dimitrijević, M.S., Moreau N., 2012. Stark-B database, [online]. Available: http://stark-b.obspm.fr [May 3, 2012]. Observatory of Paris, LERMA and Astronomical Observatory of Belgrade

REFERENCES

ARTICLE	AUTHORS	SOURCE	YEAR	METHOD	ADS REFERENCE	DOI REFERENCE	OTHER REFERENCE
Stark broadening of spectral lines of multicharged ions of astrophysical interest. II: Si IV lines.	Dimitrijević M.S., Sahal-Bréchot S., Bommier V.	A&AS, Vol.89, p.591-598	1991	SCP	http://cdsads.u-strasbg.fr/abs/1991A%26AS...89..591D	Not available	Not available
Stark broadening parameters tables for spectral lines of multicharged ions of astrophysical interest. II: Si IV lines	Dimitrijević M.S., Sahal-Bréchot S., Bommier V.	Bull. Obs. Astron. Belgrade, Vol.144, p.81-99	1991	SCP	http://cdsads.u-strasbg.fr/abs/1991BOBeo.144...81D	Not available	Not available

DATA
(DATE OF ADDITION TO THE DATABASE : 2012-03-22)

TEXT VERSION
VOTABLE VERSION

								ELECTRON			HYDROGEN II			HELIUM II			
N (CM-3)	LOWER LEVEL	UPPER LEVEL	MULTIPLET	WAVELENGTH (Å)	C (Å/CM-3)	T (K)	A	*W	W (Å)	*D	D (Å)	*W	W (Å)	*D	D (Å)	*W	W (Å)	*D	D (Å)
1.000e+17	2p6.3s 2S	2p6.3p 2Po	1	1396.7	1.400e+20	20000			1.760e-2		5.190e-5		9.070e-5		-3.040e-5		1.590e-4		-3.040e-5
1.000e+17	2p6.3s 2S	2p6.3p 2Po	1	1396.7	1.400e+20	50000			1.120e-2		-1.740e-4		2.720e-4		-7.840e-5		3.890e-4		-7.610e-5
1.000e+17	2p6.3s 2S	2p6.3p 2Po	1	1396.7	1.400e+20	80000			8.990e-3		-1.660e-4		4.040e-4		-1.180e-4		5.070e-4		-1.100e-4
1.000e+17	2p6.3s 2S	2p6.3p 2Po	1	1396.7	1.400e+20	100000			8.120e-3		-1.190e-4		4.620e-4		-1.390e-4		5.710e-4		-1.290e-4
1.000e+17	2p6.3s 2S	2p6.3p 2Po	1	1396.7	1.400e+20	150000			6.840e-3		-1.620e-4		5.710e-4		-1.810e-4		6.700e-4		-1.570e-4
1.000e+17	2p6.3s 2S	2p6.3p 2Po	1	1396.7	1.400e+20	200000			6.120e-3		-1.870e-4		6.490e-4		-2.050e-4		7.070e-4		-1.810e-4

Reset page

C IV, electron Hydrogen II Helium II

1.00e+14

or enter a value: OK

1s2.6f 2Fo -> 1s2.7g 2G

or a wavelength interval

min (Å): max (Å): OK

All

When using these data, please refer to the original papers and to this database as :

Sahal-Bréchot, S., Dimitrijević, M.S., Moreau N., 2012. Stark-B database, [online]. Available: http://stark-b.obspm.fr [Jun 4, 2012]. Observatory of Paris, LERMA and Astronomical Observatory of Belgrade

REFERENCES

ARTICLE	AUTHORS	SOURCE	YEAR	METHOD	ADS REFERENCE	DOI REFERENCE	OTHER REFERENCE
Stark broadening of spectral lines of multicharged ions of astrophysical interest. I: C IV lines	Dimitrijević M.S., Sahal-Bréchot S., Bommier V.	A&AS, Vol.89, p.581-590	1991	SCP	http://cdsads.u-strasbg.fr/abs/1991A%26AS...89..581D	Not available	Not available
Stark broadening parameters tables for spectral lines of multicharged ions of astrophysical interest. I: C IV lines	Dimitrijević M.S., Sahal-Bréchot S., Bommier V.	Bull. Obs. Astron. Belgrade, Vol.144, p.65-79	1991	SCP	http://cdsads.u-strasbg.fr/abs/1991BOBeo.144...65D	Not available	Not available
Stark broadening parameters tables for large quantum number. C IV UV lines of interest for Extreme Ultraviolet Explorer Mission	Dimitrijević M.S., Sahal-Bréchot S.	Bull. Obs. Astron. Belgrade, Vol.146, p.105-114	1992	SCP	http://cdsads.u-strasbg.fr/abs/1992BABel.146..105D	Not available	http://babel.aob.rs/
Influence of different collisional processes on the Stark broadening within spectral series : the O VI case	Dimitrijević M.S.	Bull. Obs. Astron. Belgrade, Vol.146, p.115-119	1992	SCP	http://cdsads.u-strasbg.fr/abs/1992BABel.146..115D	Not available	http://babel.aob.rs/

DATA
(DATE OF ADDITION TO THE DATABASE : 2012-05-23)

TEXT VERSION
VOTABLE VERSION

								ELECTRON			HYDROGEN II				HELIUM II			
N (CM-3)	LOWER LEVEL	UPPER LEVEL	MULTIPLET	WAVELENGTH (Å)	C (Å/CM-3)	T (K)	A	*W	W (Å)	*D	D (Å)	*W	W (Å)	*D	D (Å)	*W	W (Å)	*D	D (Å)
1.000e+14	1s2.6f 2Fo	1s2.7g 2G		7726.7	2.600e+13	20000			8.710e-2		-1.320e-3	*		*		*		*	
1.000e+14	1s2.6f 2Fo	1s2.7g 2G		7726.7	2.600e+13	50000			6.540e-2		-6.090e-4	*	2.610e-1	*	2.010e-1	*	1.660e-1		-1.830e-1
1.000e+14	1s2.6f 2Fo	1s2.7g 2G		7726.7	2.600e+13	100000			5.150e-2		2.620e-4	*	3.150e-1	*	2.550e-1	*	2.110e-1	*	1.910e-1
1.000e+14	1s2.6f 2Fo	1s2.7g 2G		7726.7	2.600e+13	200000			3.990e-2		1.510e-4	*	2.960e-1	*	2.810e-1	*	3.210e-1	*	2.120e-1

MOLAT
l'Observatoire de Paris
Astronomical Observatory, Belgrade

STARK-B: Next steps

- Insertion of MSE data
- Insertion of little “applets” on line for users:
  - ✓ Fitting along temperatures (for astrophysics)
  - ✓ extrapolation or interpolation (regularities and systematic trends)
    - along principal quantum numbers,
    - charge of the radiating ions (isoelectronic sequences),
    - homologous ions,
    - charge of the ion collider
- Future:
  - ✓ SCP code on line: STARK-C project
  - ✓ Insertion of quantum data in intermediate coupling: SST + DW

    especially adapted to highly charged ions and resonance lines *(Sahal-Bréchot with Elabidi & Ben Nessib (2004 and after, an also with Dubau and Cornille 2007 and after)*

VAMDC project

Virtual Atomic and Molecular Data Center

- European project FP7 "Research Infrastructures
  summer 2009 - end of 2012
- Interoperable e-Infrastructure
  for exchange of atomic and molecular data
- 15 administrative partners: 24 teams
  - from 6 European Union member states,
  - Serbia, Russian Federation and Venezuela
- strong coupling
  - to the users (astrochemistry, atmospheric physics, plasmas)
  - scientists and engineers from the ICT community *(Information and Communication Technologies*)
    used to deal with deploying interoperable e-infrastructure
    e.g. Europlanet IDIS
    IVOA (*International Virtual Observatory Alliance*)
    Members: Euro-VO , AstroGrid…

VAMDC
Virtual Atomic and Molecular Data Centre
http://www.vamdc.org
http://www.vamdc.eu

VAMDC: portal user

http://portal.vamdc.org/vamdc_portal_test/home.seam

Name	Description	Maintainer	Status
Cologne Database for Molecular Spectroscopy: VAMDC-TAP service	The Cologne Database for Molecular Spectroscopy (CDMS) contains a catalog of radio frequency and microwave to far-infrared spectral lines of atomic and molecular species that (may) occur in the interstellar or circumstellar medium or in planetary atmospheres. The catalog is continuously updated.	endres@ph1.uni-koeln.de	OK
ICB Dijon Methane	Calculated line lists for methane (12CH4, 13CH4 and 12CH3D). The data on methane contain the vibration-rotation energy levels, line positions and line intensities in the range from 0 to 6200 cm-1.	Christian.Wenger@u-bourgogne.fr	OK
VALD (atoms)	The Vienna Atomic Line Database (VALD) is a collection of atomic line parameters (wavelengths, transition energies and quantum numbers, oscillator strengths, Lande factors, radiative and collisional broadening). This resource is the VAMDC-TAP representation of the atomic data in VALD3.	thomas.marquart@fysast.uu.se	OK
Carbon Dioxide Spectroscopic Databank (VAMDC-TAP)	The current version of CDSD-296 consists of 7 most abundant in the Earth's atmosphere isotopic species of the carbon dixoide molecule: 626, 636, 628, 627, 638, 637, 828, covers 5.9 - 12784.1 cm-1 spectral range and contains 419610 lines.	vip@lts.iao.ru	OK
BASECOL: VAMDC-TAP interface	This database, called BASECOL is devoted to collisional ro-vibrational excitation of molecules by colliders such as atom, ion, molecule or electron. It is supervised by an international working group of molecular physicists and astrophysicits involved in the calculations and use of ro-vibrational cross-sections, in order to ensure the continuity and the quality of the database.	misha@doronin.org	OK
TOPbase : VAMDC-TAP interface	TOPbase lists LS-coupling energy levels, gf-values and photoionization cross sections for astrophysically abundant ions (Z=1,14; Z=16; Z=18; Z=20; Z=26) computed in the Opacity Project.	nicolas.moreau@obspm.fr - franck.delahaye@obspm.fr	OK
Theoretical spectral database of polycyclic aromatic hydrocarbons	The Cagliari/Toulouse PAH database is a collection of theoretical spectroscopic data about Polycyclic Aromatic Hydrocarbons and carbon clusters. It provides basic geometric characteristics, energetics, harmonic analyses and electronic photoabsorption data. It is maintained by the Astrochemistry group at INAF-Observatory of Cagliari and the Institut de Recherche en Astrophysique et PlanÃ©tologie in Toulouse.	gmulas@oa-cagliari.inaf.it	OK
Chianti	Chianti consists of a critically evaluated set of up-to-date atomic data, together with user-friendly programs written in Interactive Data Language (IDL), to analyse the spectra from astrophysical plasmas. The VAMDC interface presents just the data from the Chianti-v7 release.	gtr@ast.cam.ac.uk	OK
TIPbase : VAMDC-TAP interface	TIPbase lists fine-structure levels, A-values, collision strengths and effective collision strengths for astrophysically abundant ions, mainly from the Fe isonuclear sequence computed in the Iron Project.	nicolas.moreau@obspm.fr - franck.delahaye@obspm.fr	OK
GSMA Reims S&MPO	Calculated line lists for ozone (16O3, 16O18O16O and 18O3). The data on methane contain the vibration-rotation energy levels, line positions and line strengths in the range from 0 to 8000 cm-1.	ylb@iao.ru, vladimir.tyuterev@univ-reims.fr	OK
GSMA Reims Ethylene	Calculated data of ethylene (12C2H4). The data on ethylene contain the vibration-rotation energy levels, line positions and line intensities in the range from 500 to 7500 cm-1	ludovic.daumont@univ-reims.fr, maud.rotger@univ-reims.fr	OK
TAP-XSAMS for GhoSST database		bernard.schmitt@obs.ujf-grenoble.fr	OK
Lund laboratory spectroscopy database	Experimental data for transitions and lifetimes	hampus@astro.lu.se	OK
Stark-b	Database for "Stark" broadening of isolated lines of atoms and ions in the impact approximation	sylvie.sahal-brechot@obspm.fr	OK
Spectr-W3	The information accumulated in the SPECTR-W3 ADB contains over 450,000 records and includes factual experimental and theoretical data on ionization potentials, energy levels, wavelengths, radiation transition probabilities, oscillator strengths, and (optionally) the parameters of analytical approximations of electron-collisional cross-sections and rates for atoms and ions. Those data were extracted from publications in physical journals, proceedings of the related conferences, special-purpose publications on atomic data, and provided directly by authors. The information is supplied with references to the original sources and comments, elucidating the details of experimental measurements or calculations, where necessary and available. To date, the SPECTR-W3 ADB is the largest factual database in the world containing the information on spectral properties of multicharged ions.	p_a_loboda@mail.ru	OK
Water internet Accessible Distributed Information System	Database containing information on water spectras, notably data on H216O, HDO, D2O, H217O and H218O.	faz@iao.ru	OK
HITRAN-UCL resource	The HITRAN database - truncated version for beta testing, from http://www.cfa.harvard.edu/HITRAN/	christian.hill@ucl.ac.uk	OK
VALD sub-set in Moscow (obs)	The part of Vienna Atomic Line Database (VALD) with accurate wavelength and energy levels. It also provides laboratory and calculated transition probabilities, Lande factors and broadening parameters. It is used for line identification and spectral synthesis.	pakhomov@inasan.ru	OK

Query

Home VAMDC databases Query Saved queries Help Login Register

Query by...
Species
Processes
Environment
Advanced

Find data Save query

Legend
available, can answer
available, don't support query
unsupported keyword

- Cologne Database for Molecular Spectroscopy: VAMDC-TAP service
- ICB Dijon Methane
- VALD (atoms)
- Carbon Dioxide Spectroscopic Databank (VAMDC-TAP)
- BASECOL: VAMDC-TAP interface
- TOPbase : VAMDC-TAP interface
- Theoretical spectral database of polycyclic aromatic hydrocarbons
- Chianti
- TIPbase : VAMDC-TAP interface
- GSMA Reims S&MPO
- GSMA Reims Ethylene
- TAP-XSAMS for GhoSST database
- Lund laboratory spectroscopy database
- Stark-b
- Spectr-W3
- Water internet Accessible Distributed Information System
- HITRAN-UCL resource
- VALD sub-set in Moscow (obs)

Query by species

Home VAMDC databases Query Saved queries Help Login Register

Query by...
- Species
  - Atom
  - Molecule
  - Particle
- Processes
- Environment
- Advanced

Find data | Save query

Legend

available, can answer
available, don't support query
unsupported keyword

- Cologne Database for Molecular Spectroscopy: VAMDC-TAP service
- ICB Dijon Methane
- VALD (atoms)
- Carbon Dioxide Spectroscopic Databank (VAMDC-TAP)
- BASECOL: VAMDC-TAP interface
- TOPbase : VAMDC-TAP interface
- Theoretical spectral database of polycyclic aromatic hydrocarbons
- Chianti
- TIPbase : VAMDC-TAP interface
- GSMA Reims S&MPO
- GSMA Reims Ethylene
- TAP-XSAMS for GhoSST database
- Lund laboratory spectroscopy database
- Stark-b
- Spectr-W3
- Water internet Accessible Distributed Information System
- HITRAN-UCL resource
- VALD sub-set in Moscow (obs)

Query by species: atom

Home VAMDC databases Query Saved queries Help Login Register

Query by...
- Species
  - Atom
  - Molecule
  - Particle
- Processes
- Environment
- Advanced

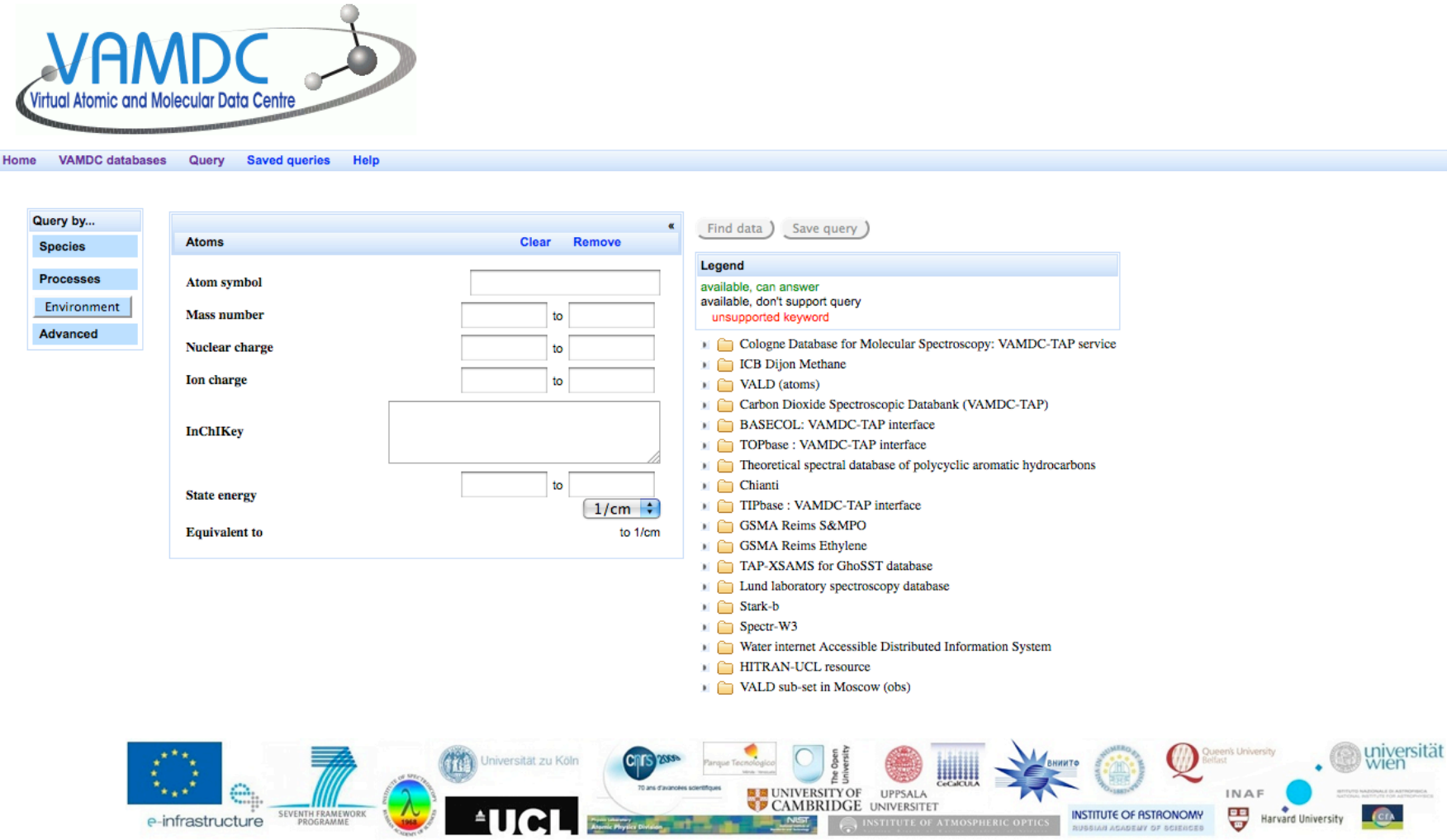
Find data | Save query

Legend

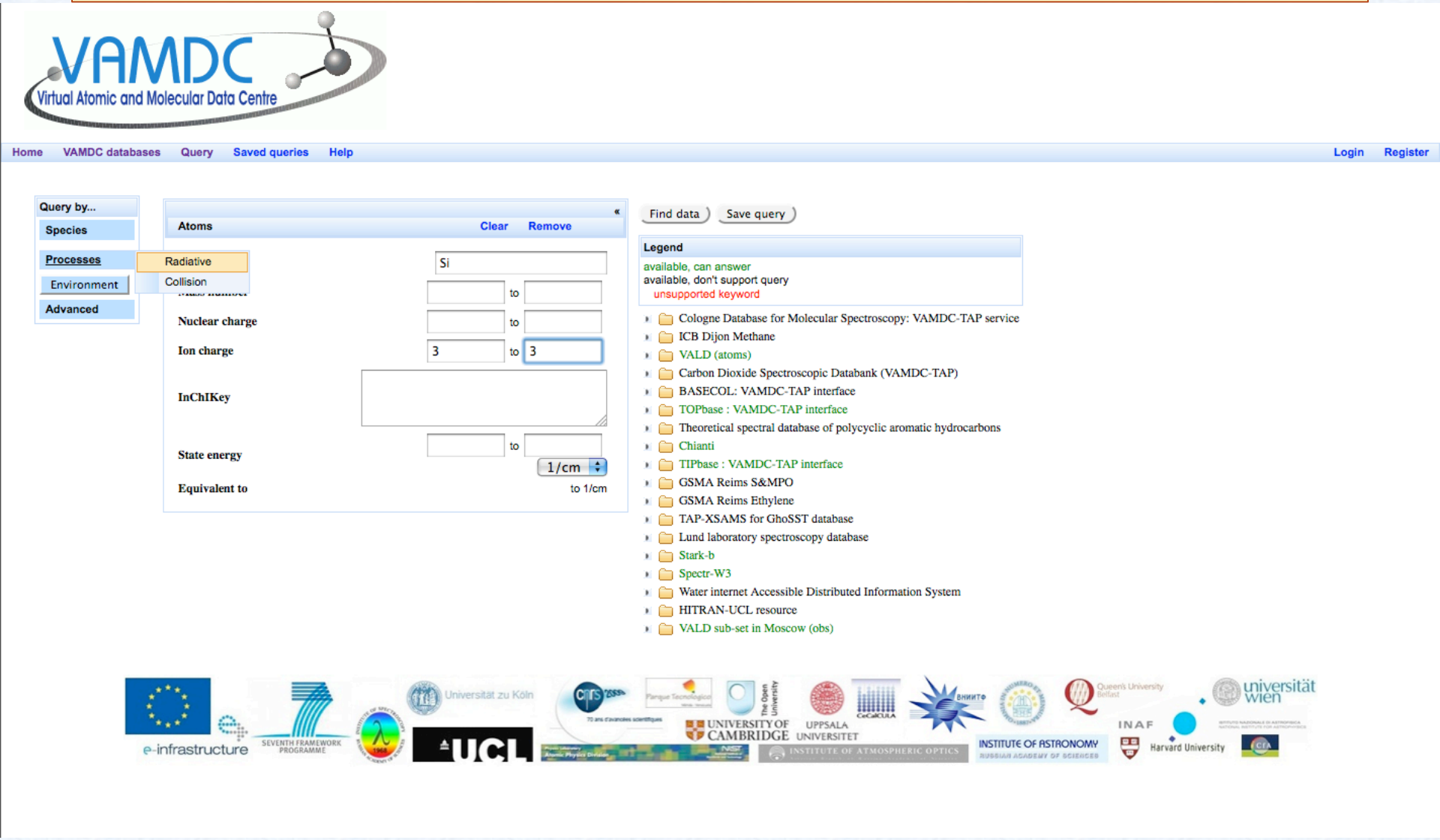
available, can answer
available, don't support query
unsupported keyword

- Cologne Database for Molecular Spectroscopy: VAMDC-TAP service
- ICB Dijon Methane
- VALD (atoms)
- Carbon Dioxide Spectroscopic Databank (VAMDC-TAP)
- BASECOL: VAMDC-TAP interface
- TOPbase : VAMDC-TAP interface
- Theoretical spectral database of polycyclic aromatic hydrocarbons
- Chianti
- TIPbase : VAMDC-TAP interface
- GSMA Reims S&MPO
- GSMA Reims Ethylene
- TAP-XSAMS for GhoSST database
- Lund laboratory spectroscopy database
- Stark-b
- Spectr-W3
- Water internet Accessible Distributed Information System
- HITRAN-UCL resource
- VALD sub-set in Moscow (obs)

Query by species: atom (following)



Query by species: atom, processes



Query by...
Species
Processes
Environment
Advanced

Atoms Clear Remove

Atom symbol	Si		
Mass number		to	
Nuclear charge		to	
Ion charge	3	to	3
InChIKey			
State energy		to	
	1/cm		
Equivalent to	to 1/cm		

Environment Clear Remove

Temperature	10000	to 50000	K
Equivalent to	10000.0 to 50000.0 K		
Pressure		to	Pa
Equivalent to	to Pa		
Number Density	1E17	to 1E17	1/cm3

Radiative Clear Remove

Wavelength	2000	to 3000	A
Equivalent Wavelength	2000.0 to 3000.0A		
Upper state energy		to	
	1/cm		
Equivalent to	to 1/cm		
Lower state energy		to	
	1/cm		
Equivalent to	to 1/cm		
Probability, A		to	1/s

Find data Save query

Legend
available, can answer
available, don't support query
unsupported keyword

- Cologne Database for Molecular Spectroscopy: VAMDC-TAP service
- ICB Dijon Methane
- VALD (atoms)
- Carbon Dioxide Spectroscopic Databank (VAMDC-TAP)
- BASECOL: VAMDC-TAP interface
- TOPbase : VAMDC-TAP interface
- Theoretical spectral database of polycyclic aromatic hydrocarbons
- Chianti
- TIPbase : VAMDC-TAP interface
- GSMA Reims S&MPO
- GSMA Reims Ethylene
- TAP-XSAMS for GhoSST database
- Lund laboratory spectroscopy database
- Stark-b
- Spectr-W3
- Water internet Accessible Distributed Information System
- HITRAN-UCL resource
- VALD sub-set in Moscow (obs)

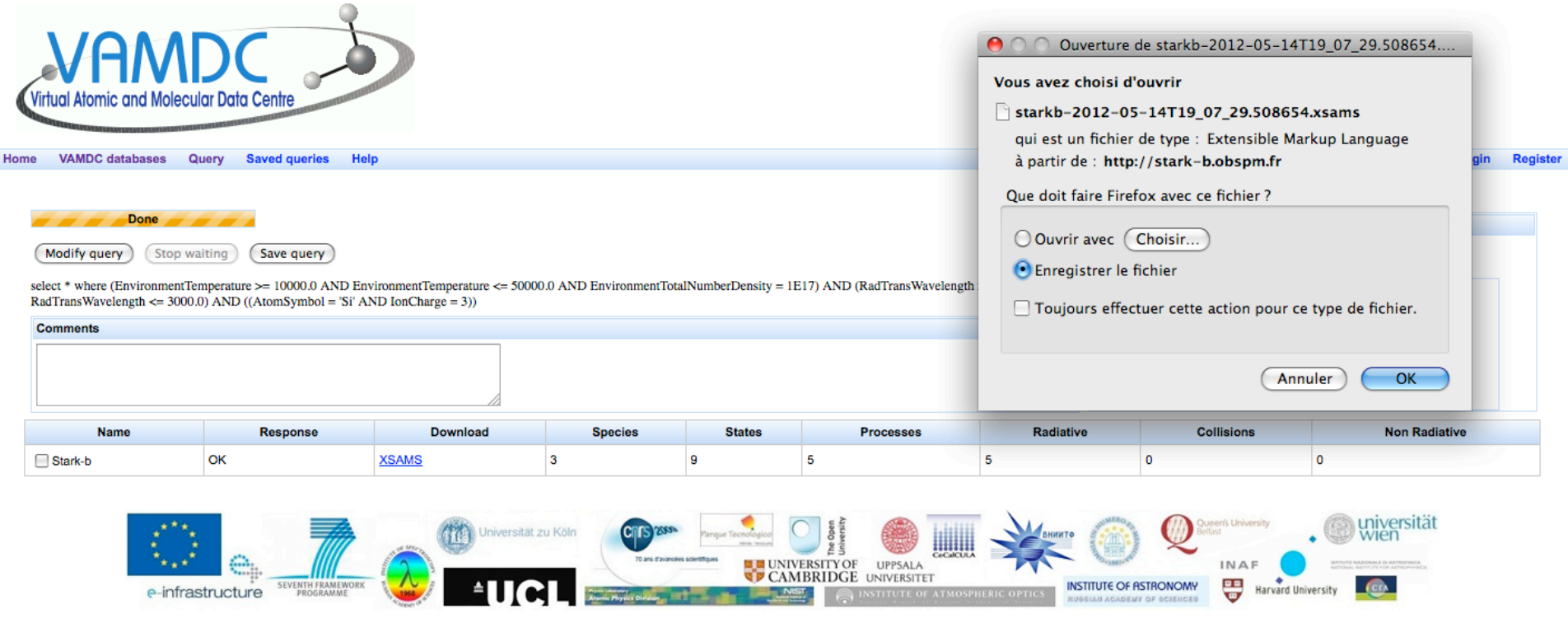
Query : result, click on download

Xsams: XML Schema for Atoms, Molecules and Solids
XML: Extensible Markup Language
Xsams2SME: converts XML document into the CSV-format wanted by Spectroscopy Made Easy (SME)

XSAMS
XML Schema for Atoms, Molecules, and Solids

IAEA
International Atomic Energy Agency
Atoms for Peace

Result of the download: code in XML language for users



Xsams
XML Schema for Atoms, Molecules and Solids

Home VAMDC databases Query Saved queries | Info Known issues Login Register

Done

Modify query Stop waiting Save query

select * where (EnvironmentTemperature >= 10000.0 AND EnvironmentTemperature <= 50000.0 AND EnvironmentTotalNumberDensity = 1E17) AND (RadTransWavelength >= 2000.0 AND RadTransWavelength <= 3000.0) AND ((AtomSymbol = 'Si' AND IonCharge = 3))

Comments

XSAMS processors

BibTeX from XSAMS
Table views of XSAMS
Xsams2SME

Process

General views of data in XSAMS format. The display is tabular and textual. Initial display is a list of states, with links to details of each state. An alternate display of radiative transitions is available.

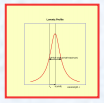
Name	Response	Download	Species	States	Processes	Radiative	Collisions	Non Radiative
Stark-b	OK	XSAMS	3	9	5	5	0	0

XSAMS processors

BibTeX from XSAMS
Table views of XSAMS
Xsams2SME

Process Result

General views of data in XSAMS format. The display is tabular and textual. Initial display is a list of states, with links to details of each state. An alternate display of radiative transitions is available.





State-list view of XSAMS

(Switch to view of radiative transitions)

Bstarkb-2012-06-04-19-34-56

Species	State	Energy
Si^{3+}	— 2p6.4p ^2P — detail	
Si^{3+}	— 2p6.5p ^2P — detail	
Si^{3+}	— 2p6.6p ^2P — detail	
Si^{3+}	— 2p6.5s ^2S — detail	
Si^{3+}	— 2p6.4d ^2D — detail	
Si^{3+}	— 2p6.5d ^2D — detail	
Si^{3+}	— 2p6.6d ^2D — detail	
Si^{3+}	— 2p6.4f ^2F — detail	
Si^{3+}	— 2p6.5f ^2F — detail	

Line-list view of XSAMS

(Switch to view of states)

Bstarkb-2012-06-04-19-34-56

Species	λ/ν/n/E	Probability	Upper state	Lower state	Broadening
Si^{3+}	λ=2125.0 A		— 2p6.5s ^{2}S	— 2p6.4p ^{2}P	Detail
Si^{3+}	λ=2675.2 A		— 2p6.5d ^{2}D	— 2p6.4f ^{2}F	Detail
Si^{3+}	λ=2676.6 A		— 2p6.6d ^{2}D	— 2p6.5p ^{2}P	Detail
Si^{3+}	λ=2483.7 A		— 2p6.6p ^{2}P	— 2p6.5s ^{2}S	Detail
Si^{3+}	λ=2287.0 A		— 2p6.5f ^{2}F	— 2p6.4d ^{2}D	Detail

Broadening of a single radiative-transition in XSAMS

Bstarkb-2012-06-04-19-19-48

Transition ID: Pstarkb-R3057

Wavelength = 2125.0 A

Type	Temperature	Pressure	Density	Composition	Profile	Parameters	Comments
pressure	20000 K		1e+17 1/cm3	• electron	Lorentzian	gammaL = 0.205	
pressure	20000 K		1e+17 1/cm3	• Hydrogen (H^+)	Lorentzian	gammaL = 0.00567	
pressure	20000 K		1e+17 1/cm3	• Helium (He^+)	Lorentzian	gammaL = 0.00628	
pressure	50000 K		1e+17 1/cm3	• electron	Lorentzian	gammaL = 0.15	
pressure	50000 K		1e+17 1/cm3	• Hydrogen (H^+)	Lorentzian	gammaL = 0.0116	
pressure	50000 K		1e+17 1/cm3	• Helium (He^+)	Lorentzian	gammaL = 0.011	

Thank you for your attention